

Limits on Unresolved Planetary Companions to White Dwarf Remnants of 14 Intermediate-Mass Stars

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ABSTRACT

We present *Spitzer* IRAC photometry of white dwarf remnants of 14 stars with $M = 3 - 5 M_{\odot}$. We do not detect mid-infrared excess around any of our targets. By demanding a 3σ photometric excess at $4.5 \mu\text{m}$ for unresolved companions, we rule out planetary mass companions down to 5, 7, or $10 M_J$ for 13 of our targets based on the Burrows et al. (2003) substellar cooling models. Combined with previous IRAC observations of white dwarf remnants of intermediate-mass stars, we rule out $\geq 10 M_J$ companions around 40 white dwarfs and $\geq 5 M_J$ companions around 10 white dwarfs.

Subject headings: infrared: stars — planetary systems — stars: low-mass, brown dwarfs — white dwarfs

1. INTRODUCTION

Radial velocity, transit, and microlensing searches are successful in finding planets around stars less massive than $2M_{\odot}$ (see Udry & Santos 2007; Marcy et al. 2005; Gould 2009; Mazeh 2009, and references therein). However, these techniques have limitations for higher mass stars. The main problem is that intermediate- and high-mass stars are big, they are rare, and they do not have many usable lines for radial velocity studies. The short (≤ 1 Gyr) main-sequence lifetimes of intermediate-mass stars and the small radii of the remnant white dwarfs (WDs) imply that the planetary systems around massive stars can be effectively studied after the hosts have been transformed into WDs (Ignace 2001; Burleigh et al. 2002).

Compared to main-sequence stars, the significant gain in contrast makes WDs excellent targets for photometric searches for planetary companions. Massive WDs are smaller and

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fainter than average mass ($0.6M_{\odot}$) WDs, increasing the contrast for substellar companions even further. If planetary systems survive the late stages of stellar evolution, massive planets should be detectable around WDs (Burleigh et al. 2002).

Detailed dynamical simulations of the evolution of the Solar system by Duncan & Lissauer (1998) show that the giant planets are likely to remain in stable orbits for more than 10 Gyr after the Sun becomes a WD. The discovery of an $M \sin i = 3.2M_J$ planet around the sdB star V391 Pegasi by Silvotti et al. (2007) and a candidate $\geq 2.4M_J$ planet around the WD GD 66 by Mullally et al. (2009) indicate that at least some planets survive post main-sequence evolution.

There are now more than a dozen WDs known to host circumstellar debris disks (see Farihi et al. 2009; Jura et al. 2009; Kilic et al. 2008, and references therein). These disks are most likely formed by tidally disrupted asteroids (Jura 2003), and at least in one case (GD 362), the composition of the accreted matter most resembles that of the Earth/Moon system (Zuckerman et al. 2007). The atmospheric composition of other disk-polluted WDs has not yet been analyzed in such detail. Farihi et al. (2009) estimate that at least 1% to 3% of WDs with cooling ages less than about 0.5 Gyr harbor debris disks. There is indirect evidence from these debris disks that at least a few percent of WDs have remnant planetary systems (Jura 2003, 2006).

Laws et al. (2003) and Fischer & Valenti (2005) find a significant trend in the frequency of planets with increasing stellar mass. The frequency of Jovian-planets ($M > 0.8M_J$ and $a < 2.5$ AU) increases from 4% for Sun-like stars to 9% for $1.3 - 1.9M_{\odot}$ stars (Johnson et al. 2007). Of course, the analogs of all of these planets around higher mass stars would most likely be swallowed in the asymptotic giant branch (AGB) phase (Villaver & Livio 2007), and the frequency of planets around massive stars at wider separations is poorly constrained at present. In the planet formation models of Kennedy & Kenyon (2008), the fraction of stars with giant planets shows a steady increase with mass up to $3M_{\odot}$. In addition, the mass of the planets and the width of the regions where they form are predicted to increase with stellar mass. Observations of massive WDs can test these models.

Gould & Kilic (2008) identified 49 young (≤ 1 Gyr) WDs from the Palomar-Green Survey that are suitable for a mid-infrared search for planetary mass companions. Here we present *Spitzer* IRAC observations of 14 stars from that sample. Our observations are discussed in Section 2, while the spectral energy distributions and the limits on planetary companions are discussed in Section 3.

2. OBSERVATIONS

We selected our targets from the DA WDs found in the Palomar-Green Survey (Liebert et al. 2005). The advantage of this sample selection is that Liebert et al. (2005) provide T_{eff} , $\log g$, mass, cooling age, and distance estimates for all these stars. We selected the 14 brightest WDs with $M_{WD} = 0.7 - 0.9M_{\odot}$ from that sample. The physical parameters of our targets, including the initial-mass estimates derived using the initial-final mass relation of Kalirai et al. (2008), are presented in Table 1. The initial-masses of our targets range from 2.8 to $4.6 M_{\odot}$. The use of different initial-final mass relations results in a $\approx 10\%$ difference in the initial-mass estimates. We estimate the main-sequence (MS) lifetimes using the equation $t_{\text{MS}} = 10(\frac{M_{\text{MS}}}{M_{\odot}})^{-2.5}$ Gyr (Wood 1992). The total (WD + MS) ages of our targets are also presented in Table 1. Four targets are about 300 Myr old, and the remaining targets have ages ≈ 1 Gyr.

Observations reported here were obtained as part of M. Kilic’s Cycle 5 *Spitzer* Fellowship Program 474. We obtained 3.6, 4.5, 5.8, and $7.9 \mu\text{m}$ images with integration times of 30 or 100 seconds per dither, with five or nine dithers per object. We use the IRAF PHOT and IDL astrolib packages to perform aperture photometry on the individual BCD frames from the latest available IRAC pipeline reduction. Since our targets are relatively faint, we use the smallest aperture (two pixels) for which there are published aperture corrections. Following the IRAC calibration procedure, corrections for the location of the source in the array are taken into account before averaging the fluxes of each of the dithered frames at each wavelength. Channel 1 ($3.6 \mu\text{m}$) photometry is also corrected for the pixel-phase-dependence. The results from IRAF and IDL reductions are consistent within the errors. The photometric error bars are estimated from the observed scatter in the five (or nine) images corresponding to the dither positions. We also add the 3% absolute calibration error in quadrature. Finally, we divide the estimated fluxes by the color corrections for a Rayleigh-Jeans spectrum (Reach et al. 2005). These corrections are 1.0111, 1.0121, 1.0155, and 1.0337 for the 3.6, 4.5, 5.8, and $7.9 \mu\text{m}$ bands, respectively.

We present the IRAC photometry of our targets in Table 2. All of our targets are detected in 2MASS, at least in the J and H bands. All but two of our targets also have Sloan Digital Sky Survey (SDSS) photometry available. PG 0852+659 and PG 1335+701 are not covered in the SDSS Data Release 7 area. We obtained V and I band photometry of PG 1335+701 using the MDM 2.4m telescope equipped with the Echelle CCD. Two sets of 120 s exposures were obtained in each filter on UT 2009 April 5. We use observations of the standard star field PG 1323–086 (Landolt 1992) to calibrate the photometry. The photometric reductions were performed by J. Thorstensen, and kindly made available to us. PG 1335+701 has $V = 15.29$ mag and $I = 15.54$ mag. We estimate an internal accu-

racy of 0.01 mag for the photometry, but to be conservative, we adopt errors of 0.03 mag. Liebert et al. (2005) provide an estimate on the V magnitude of PG 0852+659 using photographic B –band photometry. However, the error in the photographic magnitude is larger than 0.3 mag, and therefore we do not use it in our analysis. PG 0852+659 is the only WD in our sample without accurate optical photometry.

3. RESULTS

Figure 1 presents the spectral energy distributions of four WDs with ages ≈ 300 Myr. This youth is important; young planets will be brighter than their older counterparts and therefore it should be easier to find planets around relatively young WDs. The SDSS photometric zero points differ slightly from the AB convention (Eisenstein et al. 2006). We use the corrections given in Eisenstein et al. (2006) to convert the photometry to the AB system. We calculate synthetic spectra for our targets using the model atmosphere code described by Koester (2009) and the best-fit T_{eff} and $\log g$ values from Liebert et al. (2005). We perform synthetic photometry on these WD models using the appropriate transmission curves for the SDSS, 2MASS, and IRAC filters. The resulting fluxes are then compared with the observations to find the normalization factor for the WD models. We weight fluxes by their associated error bars. The solid lines in Figure 1 present the appropriate WD model for each star, which is normalized to match the observations. Optical and near-infrared photometry helps us to constrain the predicted mid-infrared photospheric fluxes for WDs. The $4.5 \mu\text{m}$ photometry of all four WDs presented in this figure is consistent with the predicted photospheric flux from WDs within 1σ ; none of them show excess mid-infrared flux.

We use the synthetic spectra for 300 Myr old $1\text{--}25 M_J$ planets (Burrows et al. 2003) to put upper limits on companions that would escape detection. The overall agreement between the Hubeny & Burrows (2007) models and the spectral energy distributions of cold brown dwarfs in the T dwarf range suggests that the planet models that we use are appropriate for the colder planets expected around WDs. The models by A. Burrows agree reasonably well with the observed hot Jupiter emission spectra in the *Spitzer* IRAC bands (Burrows et al. 2008, H. Knutson 2009, priv. comm.). However, the predicted fluxes can vary by a factor of two due to temperature inversions in the atmosphere and water being observed in emission or absorption. These temperature inversions would make the infrared fluxes higher, and they would make it easier to find planets.

The dotted lines in Figure 1 show the combined flux from each WD plus $10, 7, 5, 2$, and $1 M_J$ companions (from top to bottom, Burrows et al. 2003), respectively. The red cross marks the 3σ upper limit of the $4.5 \mu\text{m}$ photometry. By demanding a 3σ excess from any

possible companion, we exclude $\geq 5M_J$ planets around PG 1038+634, PG 1051+274, and PG 1446+286. We exclude $\geq 10M_J$ companions around PG 1335+701. Of course, these mass limits are model-dependent.

Figure 2 presents the spectral energy distributions of the remaining 10 WDs in our sample. All of these WDs are older than about 600 Myr. We use 1 Gyr old planet models to constrain the contribution from possible companions. As in Figure 1, the solid lines show the best-fit WD models. The dotted lines show the combined flux from each WD plus 1 Gyr old 25, 20, 15, 10, 7, 5, 2, and 1 M_J companions (from top to bottom), respectively. None of the WDs in Figure 2 show significant flux excesses in the mid-infrared. We rule out planets more massive than 5, 7, or 10 M_J around these targets with 3σ confidence.

Table 3 presents the search radius and the unresolved companion limits for each star. Our two pixel search radius in IRAC images corresponds to 50 AU for the nearest, and 260 AU for the most distant WD in our sample. We expect the mass loss process to enlarge the planetary orbits by a factor of $M_{\text{MS}}/M_{\text{WD}}$ (Zuckerman & Becklin 1987). We use the initial-final mass relation of Kalirai et al. (2008) to estimate the ratio of MS and WD masses, i.e., the orbital expansion factor. The above search radii allow us to probe the WDs’ progenitors for planets to 11 AU for the nearest WD and to 55 AU for the most distant one. For the separations probed (see Table 3) at 13 of our targets, we do not find any planets more massive than $10M_J$, according to Burrows et al. (2003) models. For five of our targets, there are no planets more massive than $5M_J$ within the probed inner regions.

We extend our search to partially-resolved companions by increasing the size of the photometric aperture used in our analysis. Also, it may be possible to exclude resolved companions at larger radii by the lack of 4.5 μm point sources around our targets. One of our targets, PG 1307+354, has two nearby sources that contaminate the photometry in apertures larger than two pixels. For the remaining 13 targets, the differences between three pixel aperture photometry and two pixel aperture photometry are relatively small ($\leq 3\%$). Increasing the search radius to five or ten pixel apertures is problematic as our targets are relatively faint in the mid-infrared and there are many nearby faint sources.

The direct detection of three planets at 24, 38, and 68 AU around HR 8799 (Marois et al. 2008) and the detection of a planet at 119 AU around Fomalhaut (Kalas et al. 2008) demonstrate that giant planets exist at large separations from their host stars. The planets around HR 8799 are massive enough ($5 - 13M_J$) to be detected around some of our targets. Using a three pixel aperture and accounting for orbital expansion, we would have detected a massive planet at 24 AU around 13 of our targets, a planet at 38 AU around six of our targets, and a planet at 68 AU around only three of our targets.

4. DISCUSSION

None of the stars in our sample show mid-infrared flux excess from brown dwarfs or planetary mass companions. Our data rule out $\geq 5M_J$ companions for five WDs and $\geq 10M_J$ companions for 13 WDs.

There have been many other searches for substellar and planetary mass companions to WDs. In fact, the first candidate brown dwarf was found around a WD more than 20 years ago (Becklin & Zuckerman 1988). However, only a few more WD + brown dwarf systems have been discovered since then (Farihi & Christopher 2004; Maxted et al. 2006; Steele et al. 2009). Farihi et al. (2005) find that less than 0.5% of WDs have brown dwarf companions.

Hogan et al. (2009) performed a J -band proper motion survey of 23 WDs with Gemini, and found that $\leq 5\%$ of WDs have substellar companions. In addition, the near- and mid-infrared searches by Debes et al. (2005, 2007) and Friedrich et al. (2006) did not reveal any substellar companions to WDs. *Spitzer* IRAC currently provides the best opportunity to detect planetary mass companions around WDs. An IRAC survey of 124 nearby WDs by Mullally et al. (2007) did not find any planets. However, accurate mass and age estimates are not available for the majority of the stars in their sample, and a detailed analysis would be required to put reliable limits on possibly hidden companions.

Farihi et al. (2008) present IRAC observations of 48 WDs including 31 WDs younger than 1 Gyr (MS + WD cooling age). They use blackbody models to predict the $4.5 \mu\text{m}$ photospheric flux from their targets and search for 15% excess flux at $4.5 \mu\text{m}$. Of course, the use of WD-atmosphere models (as in our study) would be more accurate, but blackbody models are sufficient for finding 15% excesses around relatively hot WDs. They rule out $\geq 10M_J$ companions around 27 of their targets. The addition of 13 stars presented in this paper brings the total sample size to 40. None of the 40 stars in the combined sample show infrared excess due to substellar companions more massive than $10M_J$. For a limit of $5M_J$, there are a total of 10 WDs in both studies. While no planets are detected (i.e., $f = 0$), it remains conceivable that the frequency is non-zero.

Due to small sample size, we use a binomial probability distribution to derive statistical uncertainties. The probability, $P(f)$, that a survey of N stars will detect n companions, when the true frequency of companions is f is given by (Burgasser et al. 2003; McCarthy & Zuckerman 2004):

$$P_n(f) = f^n(1-f)^{N-n} \frac{N!}{(N-n)!n!}. \quad (1)$$

For $N = 40$ and $n = 0$, the probability distribution peaks at zero. Since the distribution is not symmetric about its maximum value, we report the range in frequency that delimits 34% and 68% of the integrated probability function as the mean frequency and error bars, respectively. These error bars are equivalent to 1σ limits for a Gaussian distribution. We find that the frequency of $\geq 10M_J$ companions to WDs is $1.0_{-1.0}^{+1.7}\%$. This is consistent with zero.

Figure 3 displays planet versus host star mass for all known extrasolar planets detected by the radial velocity, astrometric, transit, microlensing, and direct imaging searches as of 2009 April (based on the Extrasolar Planets Encyclopedia¹). The parameter space that is ruled out by IRAC observations of the WD remnants of intermediate-mass stars presented in this work and Farihi et al. (2008) is also shown. We note that this figure excludes seven stars with $M_{WD} > 1.1M_\odot$ from the Farihi et al. (2008) study. The initial-final mass relation of Kalirai et al. (2008) implies that these WDs are the descendants of stars with $M \geq 6.5M_\odot$. However, the initial-final mass relation is steep near the top end of the WD mass function (Williams et al. 2009), and the errors are relatively large. The initial-final mass relation of Williams et al. (2009) implies an initial-mass of $\approx 6M_\odot$ for a $1.1M_\odot$ WD. Even though these seven stars are not shown in Figure 3, they are included in our statistical analysis of 40 stars presented above.

About 60% of the Sun-like ($0.7\text{--}1.3 M_\odot$) stars with RV detected planets have planet/host star mass ratios $\geq 1\frac{M_J}{M_\odot}$ (based on the Extrasolar Planets Encyclopedia). This fraction goes down to 33% for planet/host star mass ratios greater than $2\frac{M_J}{M_\odot}$. Johnson et al. (2007) find that the Jovian-planet frequency is 4.2% for Sun-like stars and it is 8.9% for $1.3\text{--}1.9M_\odot$ stars. This trend is expected to continue for higher mass stars up to $3\text{--}4M_\odot$. Kennedy & Kenyon (2008) predict that about 20% of $3\text{--}4 M_\odot$ stars have at least one gas giant planet. They also find that the isolation masses are larger for more massive stars, and the giant planets that eventually form should also be bigger (see Fig. 2 and 3 in Kennedy & Kenyon 2008).

Assuming that the planet/host star mass ratios are similar for Sun-like and intermediate mass stars, we expect 6.6% of $3\text{--}4M_\odot$ stars to have planets with planet/host star mass ratios greater than $2\frac{M_J}{M_\odot}$. We probe this regime for 6 stars in our sample. Farihi et al. (2008) reach the same limit for 20 stars, including 10 stars with $M_{WD} > 1.1M_\odot$. All 26 stars have $M_{MS} \geq 3M_\odot$. The probability of a null detection among a sample of 26 stars when the (predicted) frequency is 6.6% is 17% (see equation 1). Thus, the null detections reported in this paper and Farihi et al. (2008) are suggestive, but do not directly conflict with previous detections of planets around intermediate-mass stars. Our sample is not large enough to put

¹<http://exoplanet.eu/>

more stringent limits on the frequency of planets around intermediate-mass stars. A detailed analysis of a larger IRAC survey of WDs like that of Mullally et al. (2007) will provide useful constraints on the frequency of massive planets around intermediate-mass stars.

If the current trend in the lack of discovery of planets around WDs continues, it would suggest one of the following:

1 - Intermediate- and high-mass stars may be inefficient in forming giant planets. For stars more massive than $3M_{\odot}$, irradiation overcomes accretion as the stars reach the main sequence relatively quickly. This pushes the snow line to 10 – 15 AU and makes formation of cores difficult (Ida & Lin 2005; Kennedy & Kenyon 2008).

2 - The planets may not survive the red giant and the AGB phases. Villaver & Livio (2007) find that planets within 5.3 AU of $5 M_{\odot}$ stars will be engulfed during the AGB phase. Of course, avoiding engulfment may not be enough for a planet to survive because a corpuscular drag resulting from its interaction with the stellar wind would decrease the planetary orbits. However, Duncan & Lissauer (1998) show that the planets will only move inward slightly due to the corpuscular drag. Villaver & Livio (2007) suggest that the maximum stellar radius is reached only for a brief period of time, and the planetary orbits would have been expanded due to mass loss by that time. Hence, they find that the tidal drag forces are negligible for orbital radii larger than the maximum stellar radius reached in the AGB phase. Since the search radii for planets around the progenitors of our sample of WDs ranges from 11 to 55 AU, the engulfment of the planets during the AGB phase cannot explain the apparent lack of high mass planets around our targets.

3 - The orbits of the planets around massive stars may become unstable during the late stages of stellar evolution. Debes & Sigurdsson (2002) suggest that planets around WDs may become unstable to close approaches with each other and the entire system may become dynamically young. Duncan & Lissauer (1998) find that the crossing times for planetary orbits depend on $(M_{\text{WD}}/M_{\text{MS}})^{3.2}$. Even though this timescale is relatively long for the Sun (~ 10 Gyr), the larger mass-reduction factor for more massive stars means that the timescale for unstable close approaches is less than 1 Gyr for planets around $M \geq 3.6M_{\odot}$ stars. In addition, since orbital semi-major axes grow during the mass loss phase, some planets will pass through resonances that could create or enhance instabilities in the system. However, detailed simulations including the effects of orbital expansion on the stability of planets around massive WDs are not currently available.

5. CONCLUSIONS

There is only one known WD that has a possible planetary companion, GD 66. The signal from an $M \sin i = 2.4M_J$ planetary companion in a 4.5 yr orbit is detected in timing measurements of this pulsating WD star (Mullally et al. 2009). However, a complete orbit has not been observed yet, and the detection remains provisional. *Spitzer* IRAC observations of GD 66 did not reveal any significant mid-infrared flux excess. Based on the substellar cooling models by Burrows et al. (2003) and looser detection criteria compared to Farihi et al. (2008) and our study, Mullally et al. (2009) place an upper limit of $5 - 7M_J$ on the mass of the companion.

So far, no other search, including the IRAC surveys by Mullally et al. (2007), Farihi et al. (2008), and this paper has been successful in finding planetary companions to WDs. The proper motion surveys in the near-infrared have ruled out resolved companions to two dozen WDs (Hogan et al. 2009). However, a proper motion survey of the majority of the WDs observed with IRAC has not been performed yet. Such a survey will be valuable for searching for resolved companions to WDs. Future studies with ground based telescopes using Adaptive Optics instruments and with the *James Webb Space Telescope (JWST)* may provide the first answers to whether Jupiter mass planets survive around WDs or not.

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Table 1. High Mass WD Targets

Object	T_{eff} (K)	$\log g$ (cm s $^{-2}$)	M_{WD} (M_{\odot})	M_{MS} (M_{\odot})	d (pc)	τ_{WD} (Myr)	$\tau_{\text{WD+MS}}$ (Myr)
PG 0852+659	19070	8.13	0.70	2.8	92	140	900
PG 1034+492	20650	8.17	0.73	3.1	80	120	720
PG 1038+634	24450	8.38	0.87	4.4	68	100	350
PG 1051+274	23100	8.37	0.86	4.3	41	110	380
PG 1108+476	12400	8.31	0.80	3.7	46	600	980
PG 1129+156	16890	8.19	0.73	3.1	36	220	830
PG 1201–001	19770	8.26	0.78	3.5	63	160	580
PG 1307+354	11180	8.15	0.70	2.8	45	630	1390
PG 1310+583	10560	8.32	0.80	3.7	21	910	1290
PG 1319+466	13880	8.19	0.73	3.1	37	400	1000
PG 1335+701	30140	8.25	0.79	3.6	108	30	420
PG 1446+286	22890	8.42	0.89	4.6	47	130	350
PG 1515+669	10320	8.40	0.86	4.3	33	1120	1390
PG 1550+183	14260	8.25	0.77	3.5	41	390	840

Table 2. IRAC Photometry of High Mass WDs

Object	3.6μm (μJy)	4.5μm (μJy)	5.8μm (μJy)	8.0μm (μJy)	Reduction Pipeline
PG 0852+659	97.7 ± 3.9	65.1 ± 3.3	44.3 ± 13.1	25.2 ± 10.6	17.2
PG 1034+492	93.5 ± 5.6	57.5 ± 3.1	40.9 ± 5.9	31.7 ± 20.3	18.5
PG 1038+634	107.3 ± 4.0	65.8 ± 2.7	50.7 ± 17.4	31.0 ± 16.6	18.5
PG 1051+274	299.1 ± 12.1	190.6 ± 8.7	126.0 ± 25.5	59.1 ± 31.2	18.5
PG 1108+476	138.4 ± 5.5	87.3 ± 4.1	69.9 ± 15.9	34.6 ± 26.4	17.2
PG 1129+156	371.8 ± 12.3	230.8 ± 11.4	154.6 ± 22.1	74.9 ± 19.1	18.5
PG 1201–001	123.1 ± 4.7	79.1 ± 3.9	50.9 ± 9.0	28.9 ± 24.8	18.5
PG 1307+354	191.8 ± 6.3	118.4 ± 6.7	95.7 ± 23.6	60.5 ± 25.8	18.5
PG 1310+583	658.5 ± 23.0	420.2 ± 19.6	287.4 ± 22.4	159.3 ± 29.6	17.2
PG 1319+466	279.9 ± 11.2	177.3 ± 6.9	110.8 ± 25.2	77.9 ± 24.3	18.5
PG 1335+701	92.1 ± 3.6	56.5 ± 3.3	30.4 ± 9.6	24.5 ± 11.0	17.2
PG 1446+286	195.1 ± 9.7	120.1 ± 6.2	99.7 ± 18.6	54.5 ± 24.7	18.5
PG 1515+669	187.7 ± 7.9	119.4 ± 7.5	94.4 ± 16.9	24.4 ± 11.9	17.2
PG 1550+183	218.3 ± 8.0	136.6 ± 5.3	96.0 ± 20.7	50.3 ± 11.7	18.5

Table 3. Limits on Unresolved Companions

Object	Search Radius [†] (AU)	Initial Separation [†] (AU)	M_{MS} (M_{\odot})	Ruled out ^{††} Companions
PG 0852+659	221	55	2.8	$> 15M_J$
PG 1034+492	192	46	3.1	$> 7M_J$
PG 1038+634	163	32	4.4	$\geq 5M_J$
PG 1051+274	98	20	4.3	$\geq 5M_J$
PG 1108+476	110	24	3.7	$\geq 5M_J$
PG 1129+156	86	20	3.1	$> 7M_J$
PG 1201–001	151	33	3.5	$\geq 10M_J$
PG 1307+354	108	27	2.8	$> 10M_J$
PG 1310+583	50	11	3.7	$\geq 10M_J$
PG 1319+466	89	21	3.1	$\geq 5M_J$
PG 1335+701	259	56	3.6	$\geq 10M_J$
PG 1446+286	113	22	4.6	$\geq 5M_J$
PG 1515+669	79	16	4.3	$\geq 7M_J$
PG 1550+183	98	22	3.5	$\geq 7M_J$

[†]For a two pixel aperture.

^{††}These limits are calculated based on the models by Burrows et al. (2003). See the discussion in Section 3.

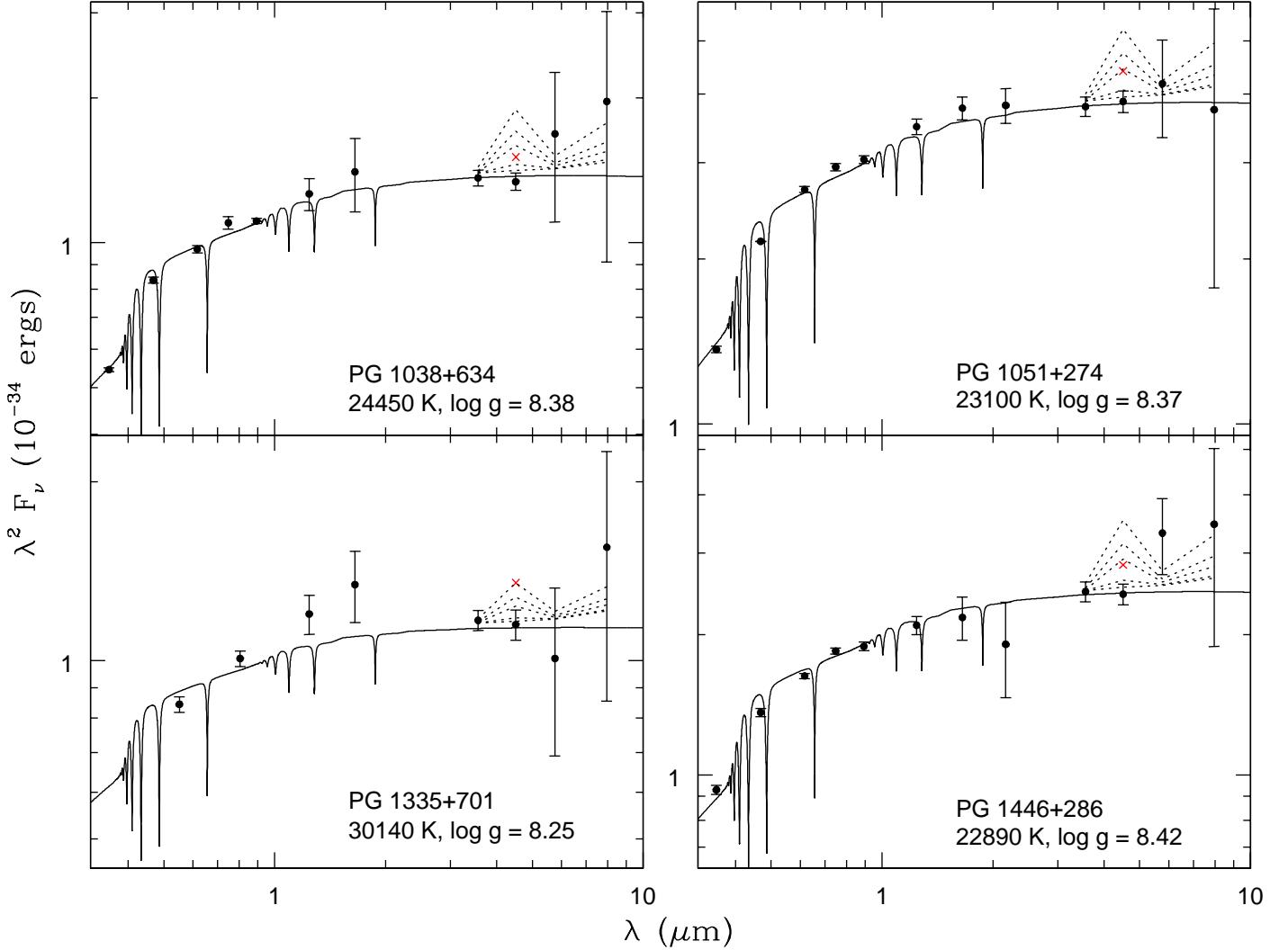


Fig. 1.— Spectral energy distributions of four 300 Myr old (total age) WDs. The expected photospheric flux from the WDs are shown as solid lines (Koester 2009). The dotted lines show the expected flux from planetary companions with $M = 10, 7, 5, 2$ and $1 M_J$ (from top to bottom), respectively. The red cross marks the 3σ upper limit of the $4.5\mu\text{m}$ photometry.

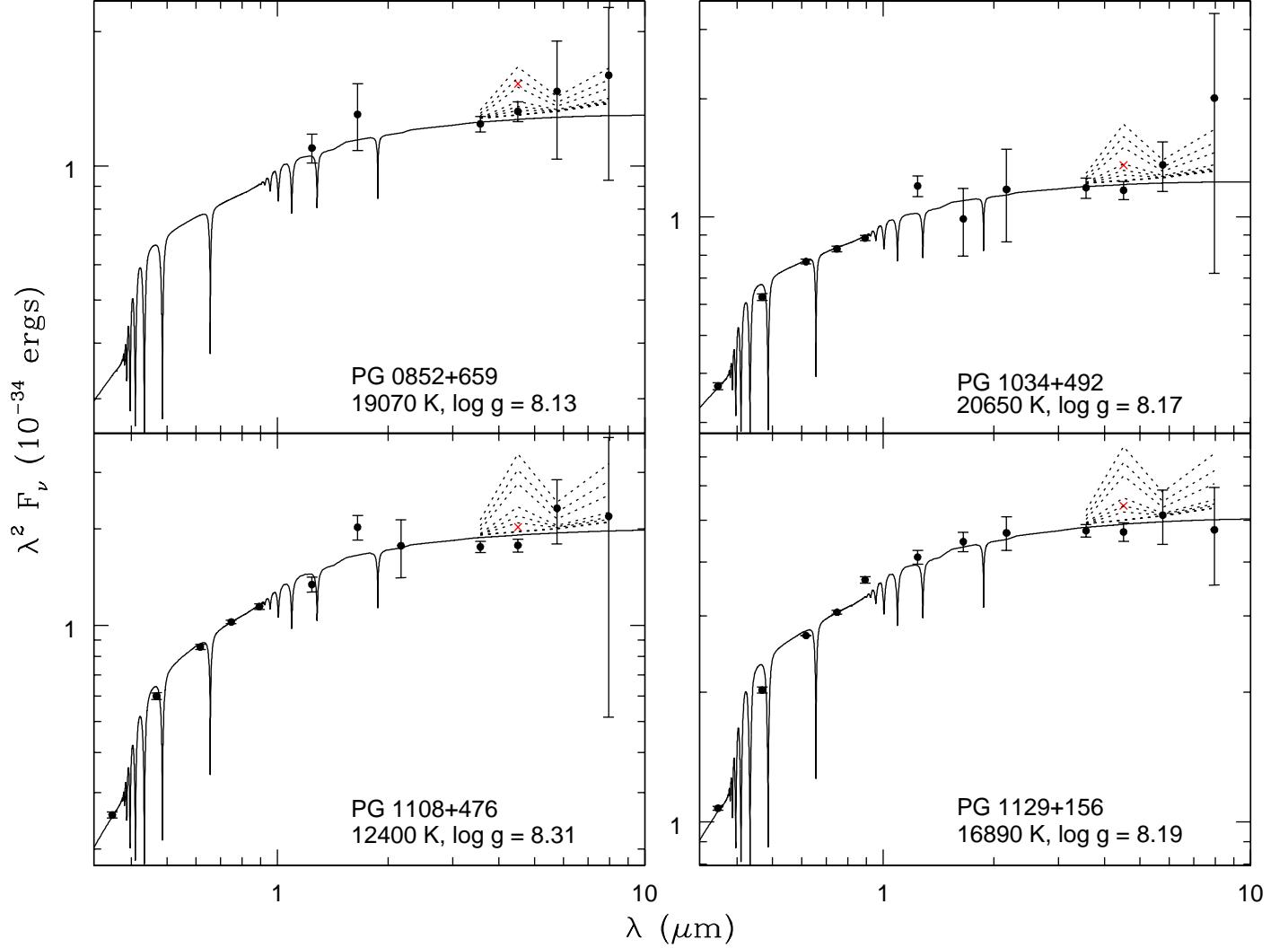


Fig. 2.— Spectral energy distributions of ≈ 1 Gyr old (total age) WDs. The expected photospheric flux from the WDs are shown as solid lines (Koester 2009). The dotted lines show the expected flux from planetary companions with $M = 25, 20, 15, 10, 7, 5, 2$ and $1 M_J$ (from top to bottom), respectively. The red cross marks the 3σ upper limit of the $4.5\mu\text{m}$ photometry.

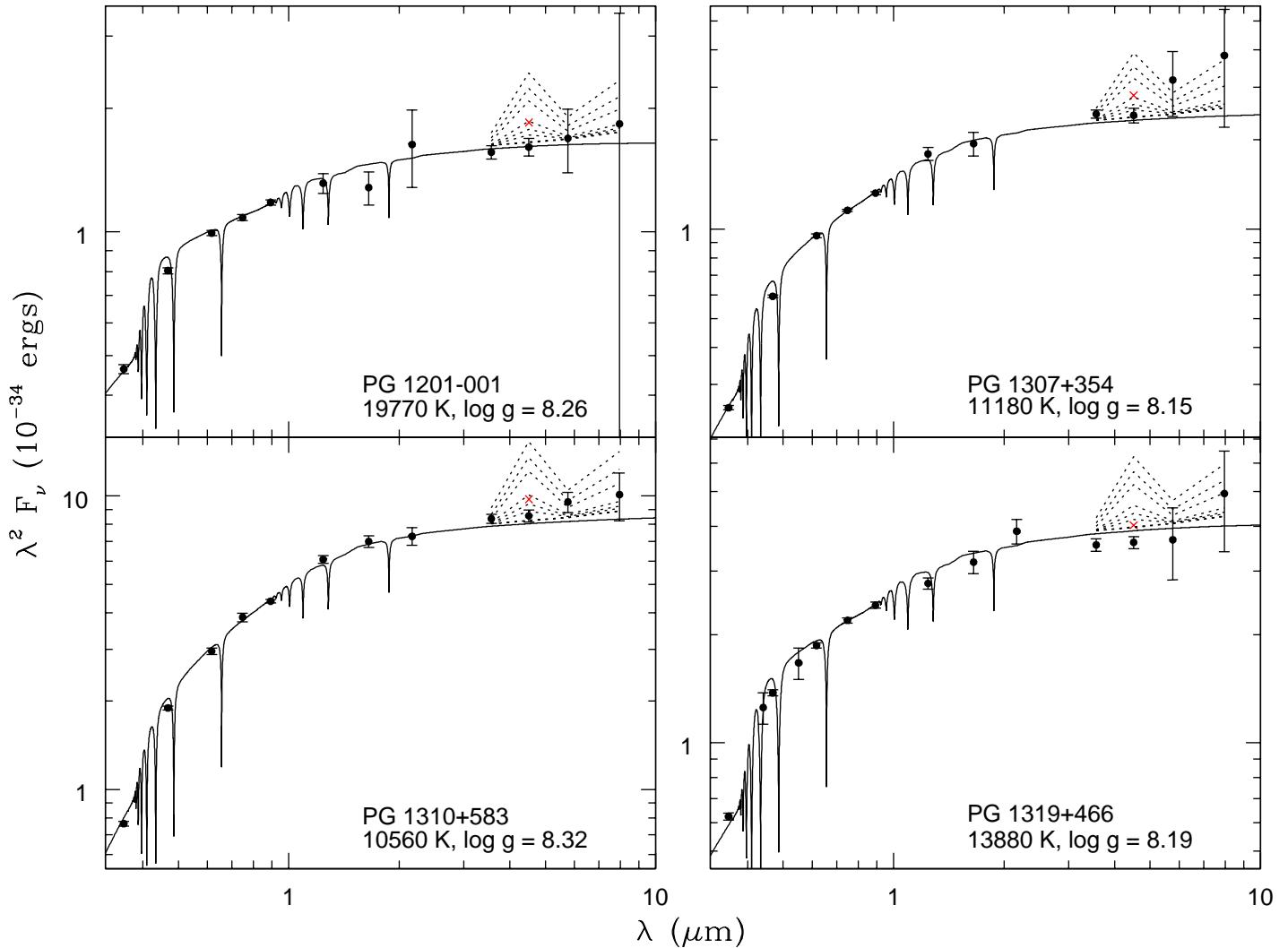


Fig. 2.— contd.

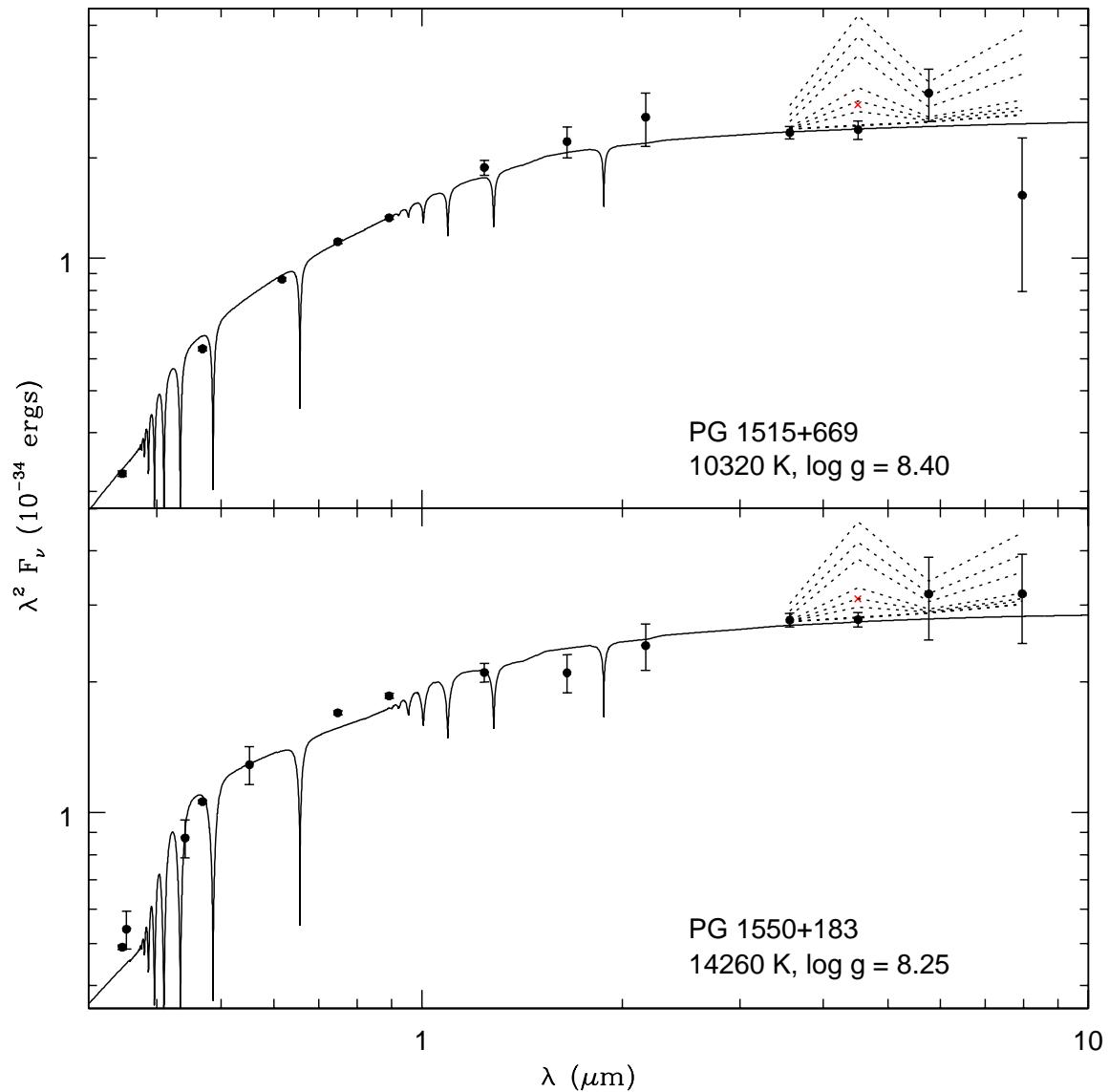


Fig. 2.— contd.

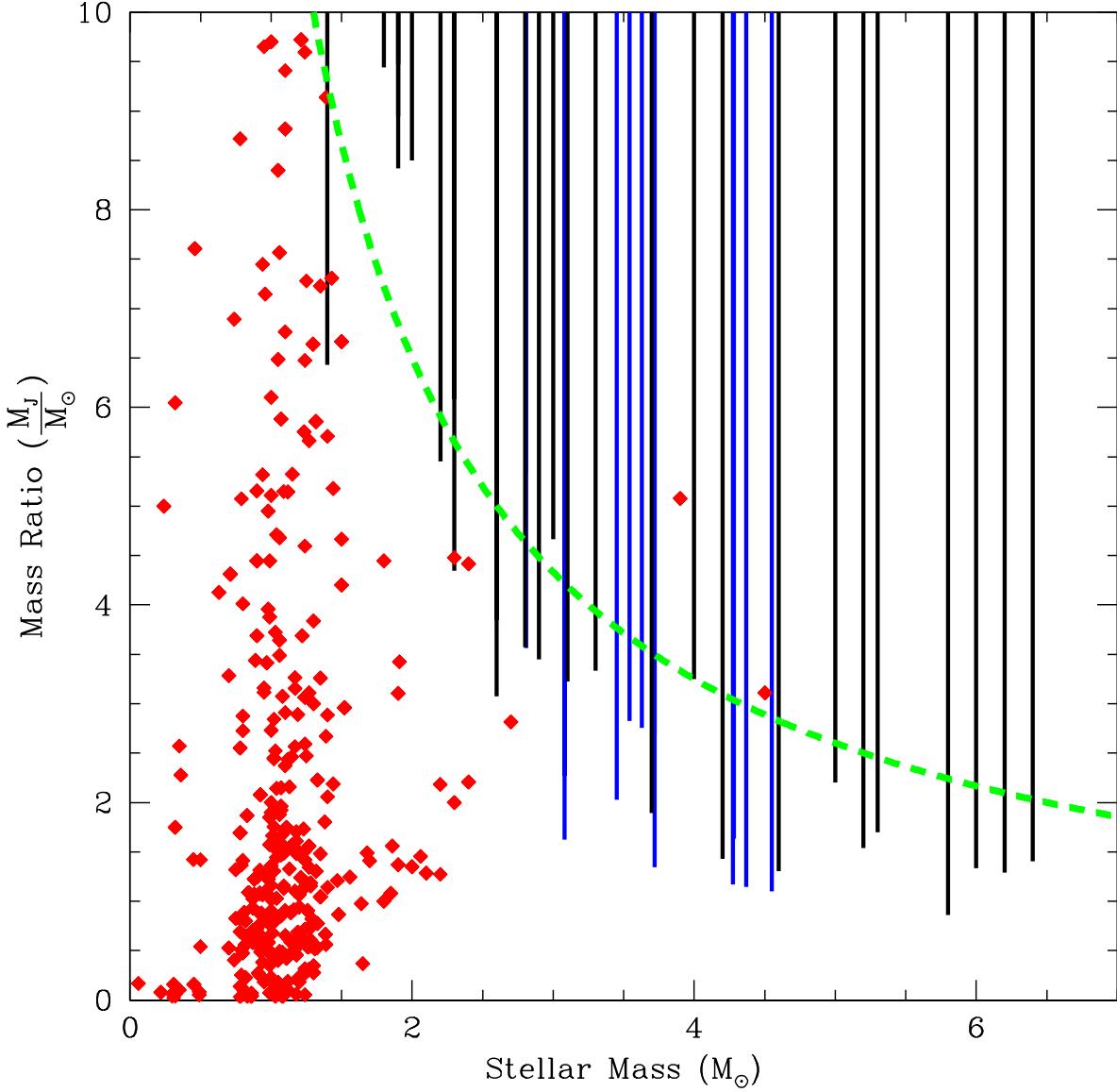


Fig. 3.— Planet versus host mass for all known extrasolar planets detected by the RV, astrometry, transits, microlensing, and direct imaging observations as of 2009 April. The dashed line marks the upper limit for planetary mass objects ($M = 13M_J$). The parameter space that is ruled out by *Spitzer* observations of WD remnants of intermediate-mass stars by Farihi et al. (2008) and this work is shown as black and blue lines, respectively. We use 3σ (12-18%), while Farihi et al. (2008) use 15% excess for planet mass upper limits.